

# Photometric variability of the Be star CoRoT<sup>★</sup>-ID 102761769

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## ABSTRACT

*Context.* Classical Be stars are rapid rotators of spectral type late O to early A and luminosity class V-III, which exhibit Balmer emission lines and often a near infrared excess originating in an equatorially concentrated circumstellar envelope, both produced by sporadic mass ejection episodes. The causes of the abnormal mass loss (the so-called Be phenomenon) are as yet unknown.

*Aims.* For the first time, we can now study in detail Be stars outside the Earth's atmosphere with sufficient temporal resolution. We investigate the variability of the Be Star CoRoT-ID 102761769 observed with the CoRoT satellite in the exoplanet field during the initial run.

*Methods.* One low-resolution spectrum of the star was obtained with the INT telescope at the Observatorio del Roque de los Muchachos. A time series analysis was performed using both cleanest and singular spectrum analysis algorithms to the CoRoT light curve. To identify the pulsation modes of the observed frequencies, we computed a set of models representative of CoRoT-ID 102761769 by varying its main physical parameters inside the uncertainties discussed.

*Results.* We found two close frequencies related to the star. They are  $2.465 \text{ c d}^{-1}$  ( $28.5 \mu\text{Hz}$ ) and  $2.441 \text{ c d}^{-1}$  ( $28.2 \mu\text{Hz}$ ). The precision to which those frequencies were found is  $0.018 \text{ c d}^{-1}$  ( $0.2 \mu\text{Hz}$ ). The projected stellar rotation was estimated to be  $120 \text{ km s}^{-1}$  from the Fourier transform of spectral lines. If CoRoT-ID 102761769 is a typical Galactic Be star it rotates near the critical velocity. The critical rotation frequency of a typical B5-6 star is about  $3.5 \text{ c d}^{-1}$  ( $40.5 \mu\text{Hz}$ ), which implies that the above frequencies are really caused by stellar pulsations rather than star's rotation.

**Key words.** stars: early-type – stars: emission-line, Be – stars: individual: CoRoT-ID 102761769 – stars: rotation – stars: oscillations

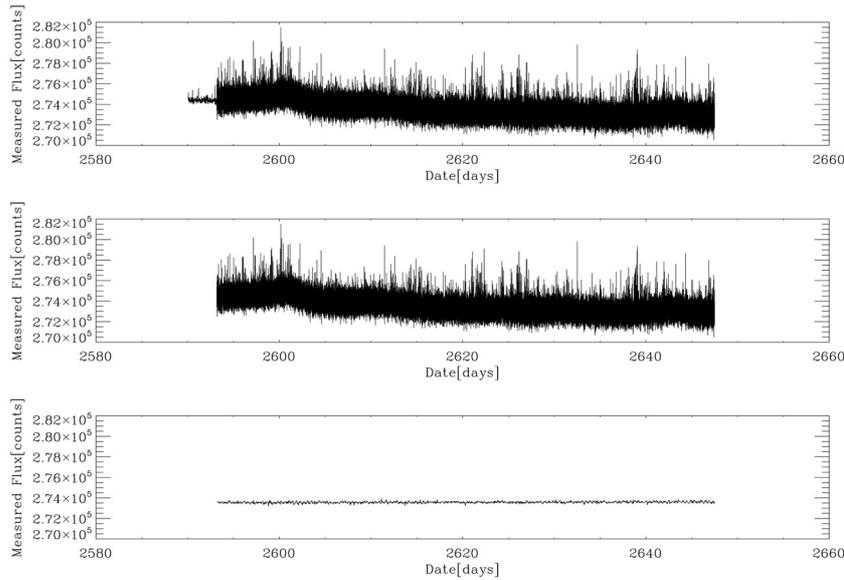
## 1. Introduction

Classical Be stars are rapid rotators of spectral type late O to early A and luminosity class V-III, which exhibit Balmer emission lines and often a near infrared excess originating in an equatorially concentrated circumstellar envelope, both produced by sporadic mass ejection episodes. The causes of the abnormal mass loss (the so-called Be phenomenon) are as yet unknown. In spite of their high  $V \sin i$  (Frémat et al. 2005), rapid rotation alone cannot explain the ejection episodes as most Be stars do not rotate at their critical rotation rates (see Porter & Rivinius 2003, for a comprehensive review of Be stars).

High-resolution, high signal-to-noise spectroscopic observations have been analyzed to demonstrate short-term variations

are rather common among early-type Be stars (Baade 1982, 1984; Porri & Stalio 1988; Floquet et al. 1992; Neiner et al. 2002). The observed line profile variability (LPV) is characterized by moving bumps traveling from blue to red across the line profile on timescales ranging from minutes to a few days. The phenomenon has also been observed in O stars and  $\delta$  Sct variables (e.g. Kennely et al. 1992), among others. Non-radial pulsations (NRP) have been proposed as an explanation of the LPV observed in hot stars (e.g. Smith 1977; Baade 1982; Vogt & Penrod 1983). NRP produce LPV thanks to the combination of the Doppler displacement of stellar surface elements with their associated temperature variations due to the compression/expansion caused by the passage of waves through the photosphere. NRP could be the additional mechanism required for a rapidly rotating B star to become a Be star, that is to trigger the Be phenomenon by means of mass ejection. Indeed, Rivinius et al. (1998) found multiperiodicity in  $\mu$  Cen and showed that a

\* The CoRoT space mission was developed and is operated by the French space agency CNES, with participation of ESA's RSSD and Science Programmes, Austria, Belgium, Brazil, Germany, and Spain.



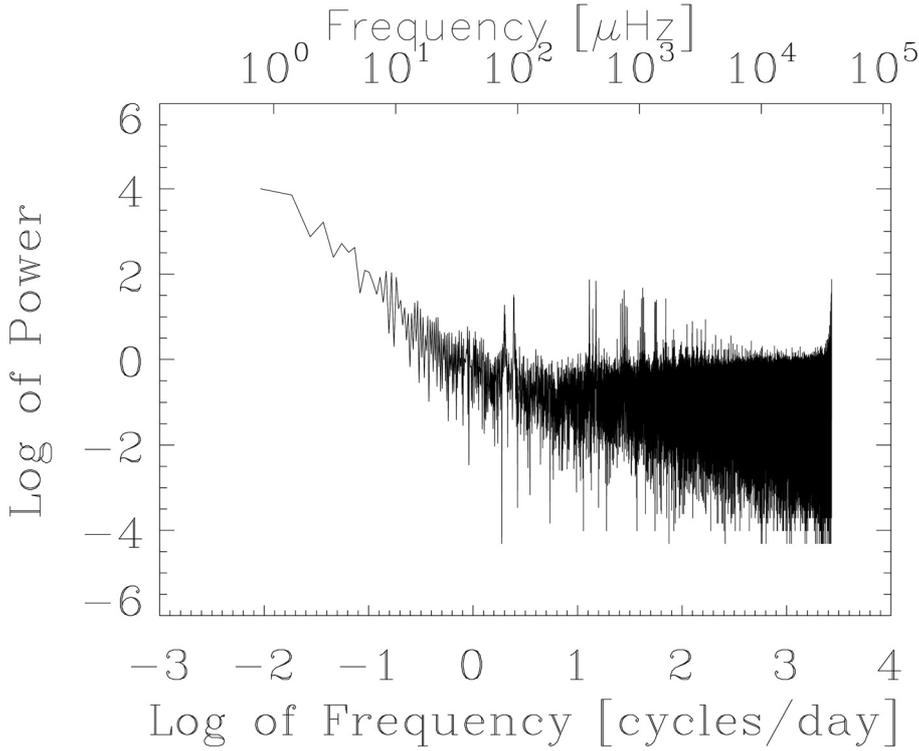
**Fig. 1.** *Top:* CoRoT-ID 102761769 star light curve. *Middle:* data after outlier removal at three  $\sigma$  level. *Bottom:* moving average every 100 points.

correlation exists between mass ejection episodes and the beating pattern of the multiperiodicity. The finding of new cases of Be stars for which beating periods of multiperiodic NRPs coincide with matter ejections would help us to confirm this model. The periodic variability of the star has been reproduced in detail by NRP modeling (Rivinius et al. 2001). Short-periodic LPV of other Be stars have also been modeled using NRP. Observations with the MOST satellite showed that multiperiodicity due to NRP is a rather common phenomenon among Be stars (e.g. Walker et al. 2005; Dziembowski et al. 2007). Nevertheless, to measure frequencies with great accuracy in stars requires both high photometric precision and high time-frequency resolution. The CoRoT (Convection, Rotation and planetary Transits) satellite opens for the first time in history the possibility of fulfilling this goal. CoRoT (Boisnard & Auvergne 2004) is an experiment of astronomy dedicated to seismology and the detection of extrasolar planet transits. It was launched on December 2006 in an inertial polar orbit at an altitude of 897 km. The instrument is fed by a  $\phi = 27$  cm telescope. Its scientific program is three-fold consisting of: (1) the seismology core program (SISMO), which concerns the seismic study of  $\sim 10$  bright ( $6 < V < 9$ ) well-selected pulsating stars per observed field. Those measurements are sampled with a cadence of 32 s, which allows very accurate frequency analysis; (2) the exoplanet core program (EXO), which concerns the search for exoplanets around  $\sim 12000$  faint ( $11.5 < V < 16$ ) stars per field. The sampling of the EXO data has a cadence of 512 s; (3) the additional program (AP), which concerns the study of a few hundreds faint ( $11.5 < V < 16$ ) variable stars per EXO field. Within the core exoplanet data, there are hot stars among which there will certainly be new pulsators, enriching the statistics of the NRP phenomenon. In this work we report on the analysis of the star CoRoT-ID 102761769. Because of a certain instability of the attitude control system (satellite jitter), the analysis of some of CoRoT stars requires a careful treatment. In the next section, we present the light curve as coming from the reduction pipeline and explain the measured instrumental noise. In Sect. 3, we discuss the time analysis of the light curve and how the information was extracted. In Sect. 4, we describe the spectroscopy ground-based observations of the star and an estimate of the rotation period. In Sect. 5, we discuss the

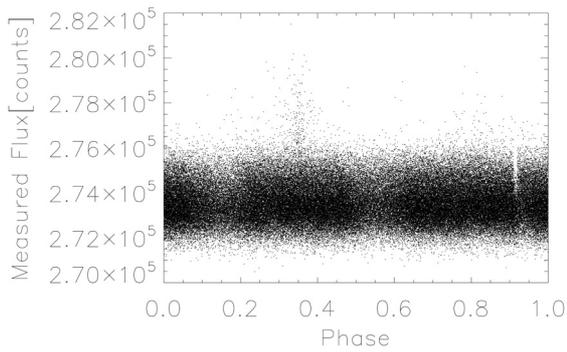
frequencies found and compare them with the critical rotation velocity. In Sect. 6, the conclusions of this work are presented.

## 2. CoRoT photometric observations

The data were taken in the exo field during the initial run (IRa01) during 54.6 days with an initial cadence of 512 s. An automatic procedure (alarm mode) triggers a higher cadence in the case of suspect detection of a planetary transit. After 3.13 days, CoRoT began to record the data of the ID 102761769 star with a cadence of 32 seconds (Fig. 1 – top). The coordinates (J2000.0) of the star are RA =  $06^{\text{h}}45^{\text{m}}02^{\text{s}}.4$  and Dec =  $-01^{\circ}50'08''.03$  and  $V = 13.288$  mag. The star is also cited in catalogues mentioning as emission-line object (Kohoutek & Wehmeyer 1999; Robertson & Jordan 1989). The light curve contains 129 305 photometric observations. Looking for oscillations, we first applied the cleanest algorithm (Foster 1995) to the data. Figure 2 shows the first cleanest power spectrum. We identified several frequencies, the most easily distinguishable ones being related to the satellite's orbit and its harmonics. The close distance from Earth causes perturbations on the satellite leading to satellite jitter (Auvergne et al. 2009). The main sources are eclipses, the variation in the Earth's gravity and magnetic field, South Atlantic anomaly, the Sun and Earth emissivity, Earth's albedo, and objects in low Earth orbit (LEO). Figure 3 shows a phase diagram of the raw data with the 6184 s CoRoT satellite orbit period. CoRoT pipeline have a flag for bad points. Some of the points were corrected by the pipeline, while others were eliminated, which is why we have missing points. The gaps in Fig. 3 were caused by points eliminated from the automatic pipeline reduction mainly by satellite jitter. The cluster of points in Fig. 3 are the remaining points that the pipeline could not correct. We eliminated 480 points greater than  $3\sigma$  from the average (0.4% of the total points). We also eliminated the first 355 observations (0.3% of the total points) taken with a larger cadence and filled the gaps with the average value (Fig. 1 – middle). The lower frequencies shown in the power spectrum in Fig. 2 are due to hot pixels (Auvergne et al. 2009). There is a prism mounted in front of the CoRoT exoplanet channel camera detectors, which



**Fig. 2.** Cleanest spectrum of the light curve (Fig. 1 – top). Most of the peaks are harmonics of the satellite orbital period due to the gaps in phase with the orbit (see Fig. 3). Among the significant frequencies, only two of them are not related to the the satellite orbital period harmonics.



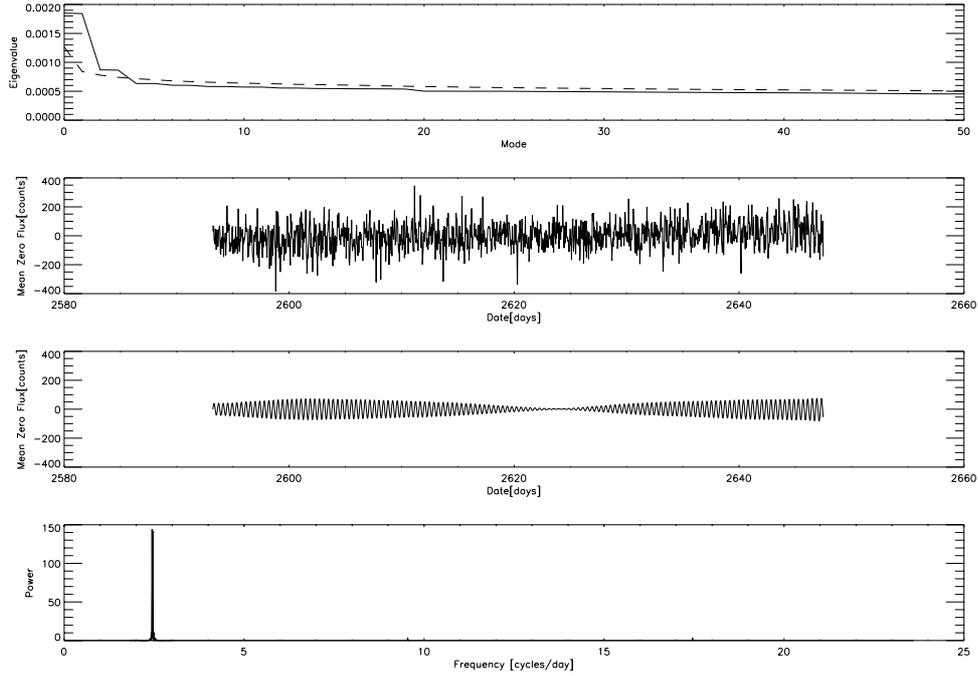
**Fig. 3.** Phase diagram of the satellite folded modulo the orbital period (102.9 min). The gaps that appear in phases 0.3 and 0.95 are data eliminated by the automatic pipeline reductions due mainly to satellite jitter.

allows CoRoT to gather 3-colour light curves. We analyzed the total sum of the light curves (the so-called white light curve) to highest quality relation in terms of signal/noise. Each color light curve falls on different pixels, what allows us to check for cosmic rays. The bump around date 2600 (Fig. 1 – middle) is only found in the red channel, and is indicative of cosmic-ray source noise. The lower frequencies were eliminated by a polynomial fit to the data and the missing points were interpolated (17 715 in total). An average of each 100 points was finally evaluated by reusing each point 100 times and then selecting 1 point every 100 points for additional analysis (Fig. 1 – bottom). The total number of points in Fig. 1 (bottom) is 1466. The regulary spaced intervals of data series were essential to the time analysis described in the next section.

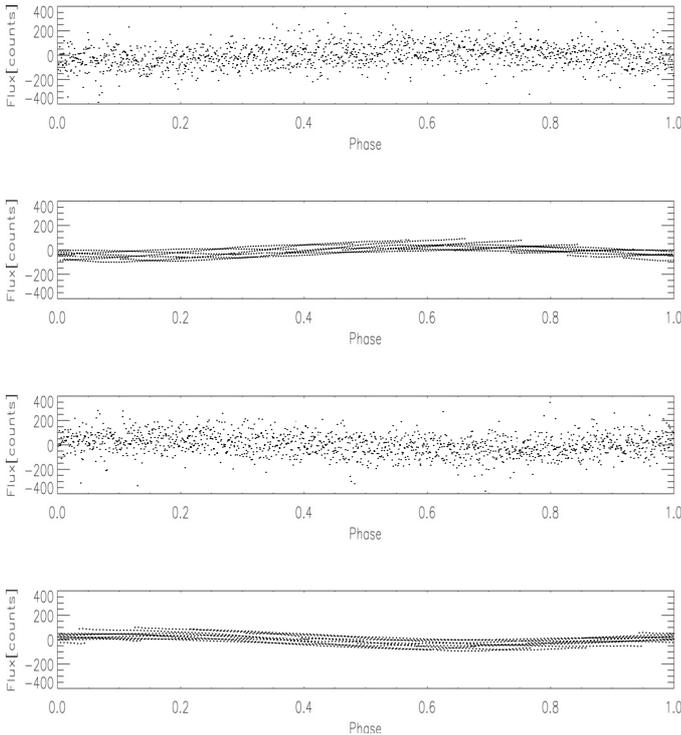
### 3. Time analysis

We applied the singular spectrum analysis (SSA) (Ghil et al. 2002) to the data in Fig. 1 – bottom. SSA is designed to extract information from noisy and evenly spaced time series. The method provides insight into the unknown or only partially

known dynamics of the underlying system that generated the series. The procedure takes an univariate record and produces a multivariate set of observations by using lagged copies of a single time record, we are able to define the coordinates of the phase space that approximate the dynamics of the system. SSA has been applied extensively to the study of climate variability as well as other areas in physical and life sciences. To create the singular spectrum we have to sort the eigenvalues of the trajectory matrix. Each eigenvalue is associated with an oscillation in the time series. The full line in Fig. 4 (top) shows the singular spectrum applied to the Be star CoRoT-ID 102761769. This plot is the strength of the sorted eigenvalues. We performed a 100 runs of Monte Carlo SSA shuffling the data (Allen & Smith 1996). The dashed line in Fig. 4 represents the  $3\sigma$  confidence limit for this analysis. This means that the first four eigenvalues, defined by our Monte Carlo analysis, almost entirely encompass the signal of our time series. We reconstructed the time series (filtered series) with this limit (3rd row of Fig. 4). For comparison, we plotted again the bottom curve of Fig. 1 on a similar scale as our filtered series in the 2nd row of Fig. 4. The reconstructed signal is a typical portrait of a beating process as expected for a Be star. Since SSA provides only the most significant eigenvalues found into data, we applied cleanest (Foster 1995) to identify the equivalent frequencies. Cleanest projects the data into an orthogonal subspace of trial functions (sines, cosines, and a constant function), while the projection provided by SSA is onto the data itself. Figure 4 (bottom) shows the cleanest spectrum of the reconstructed time series. The cleanest spectrum follow a chi-square distribution with  $r - 1$  degrees of freedom, where  $r$  is the number of trial functions (Foster 1996). The frequencies at  $2.465 \text{ c d}^{-1}$  and  $2.441 \text{ c d}^{-1}$  are more than 99.9% above the significance limit. Figure 5 shows the phase diagrams for the two frequencies found for both smoothed and filtered series. We had already found those frequencies in the original light curve. To ensure that these frequencies were not introduced by numerical manipulation, we searched for them in each step of our analysis. Figure 6 shows a zoom in the  $2.5 \text{ c d}^{-1}$  region in the cleanest

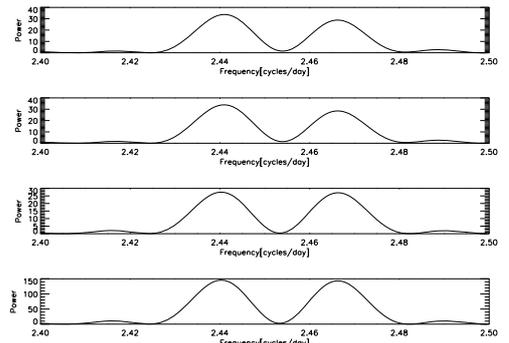


**Fig. 4.** *Top:* the full line shows the singular spectrum and the dashed line shows an average of 100 runs of Monte Carlo simulation. *2nd row:* smoothed time series (same as Fig. 1 – *bottom*). *3rd row:* reconstructed time series using only the leading 4 eigenvalues, showing a strong beating effect. *Bottom:* Fourier transform of the reconstructed series. Note that the frequencies around  $2.4 \text{ c d}^{-1}$  are by far the dominant ones.



**Fig. 5.** Phase diagrams. *From top to bottom:* smoothed data with frequency  $=2.465 \text{ c d}^{-1}$ , SSA result with frequency  $=2.465 \text{ c d}^{-1}$ , smoothed data with frequency  $=2.441 \text{ c d}^{-1}$ , SSA result with frequency  $=2.441 \text{ c d}^{-1}$ .

spectrum. From top to bottom we display, the light curve (lc), the lc without orbit perturbations, the lc with interpolated gaps, and the reconstructed time series. For all the light curves, the signal is above the significance limit. The difference between the two close frequencies is  $0.024 \text{ c d}^{-1}$  above the limit of the frequency resolution, which is  $0.018 \text{ c d}^{-1}$ . The amplitude found for

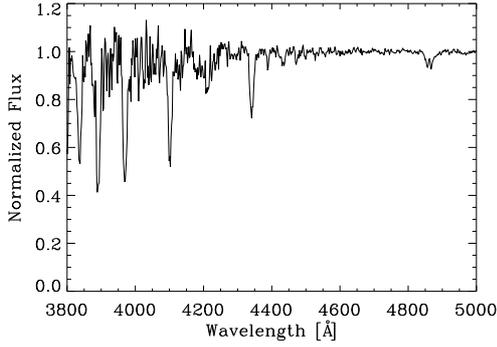


**Fig. 6.** Cleanest spectra between  $2.4$  and  $2.5 \text{ c d}^{-1}$ , *from top to bottom:* the light curve (lc), lc without orbit perturbations, lc with interpolated gaps, reconstructed time series.

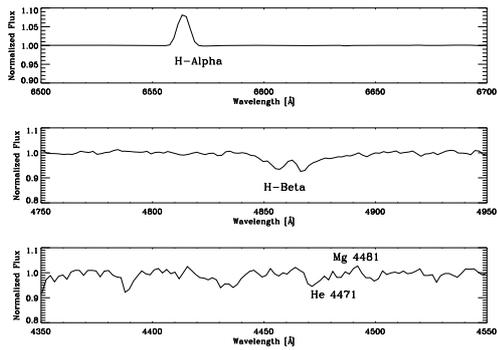
both frequencies is  $0.14 \pm 0.06 \text{ mmag}$  using the algorithm *cleanest*. The CoRoT photometric performance is  $7.7 \times 10^{-4} \text{ mag}$  in 1 h as faint as  $V = 15.5$  (Auvergne et al. 2003). Aigrain et al. (2009) evaluated the noise per exposure on 2 h timescales for the IRa01 and found the photometric performance to be close to the pre-launch specification. They found an empirical relation between 2 h noise and  $R$ -magnitude. Applying this empirical relation for the CoRoT-ID 102761769 star and taking all data analyzed (53.47 days), we found that  $\sigma_{53.47 \text{ days}} = 0.01 \text{ mmag}$ . However this light curve is particularly noisy because of satellite jitter. The standard deviation of the mean from our data is  $0.01 \text{ mmag}$ . The outliers in the same figure are still caused by the jitter that was not totally eliminated from our analysis. This is why it is difficult to see the 2 frequencies from the 2nd row of Fig. 4.

#### 4. Spectroscopy ground-based observation

A spectroscopic observation of the CoRoT-ID 102761769 star was performed at Observatorio del Roque de los Muchachos on



**Fig. 7.** Spectroscopic ground-based observation of CoRoT-ID 102761769 star.



**Fig. 8.** Close-ups of interesting spectral regions. *From top to bottom:* H $\alpha$ , H $\beta$ , He 4471 Å/Mg 4481 Å.

March 30, 2008 with the Intermediate Dispersion Spectrograph (IDS) at the 2.5 m Isaac Newton Telescope. The spectrum was taken with 300 s of exposure time using the 300 l/mm grating centred at 5460 Å and covering the region from 3337 Å to 8366 Å and resolution of 2 Å (Fig. 7). Figure 8 shows views of the H- $\alpha$ , H- $\beta$ , and He 4471 Å/Mg 4481 Å regions. The reduction was accomplished using IRAF software. The H $\alpha$  line is seen in emission and H $\beta$  is partially filled with emission. The spectral characteristics are indicative of a spectral type B5-6 IV-Ve based on the EWs of HeI lines, H $\delta$ , H $\gamma$ , and Mg II ( $\lambda$ 4481 Å). The CoRoT catalog indicates for the star an A0V spectral type and luminosity class. This classification was obtained from photometric data using a statistical method (Deleuil et al. 2009). Our spectral classification based on spectroscopic data provides a more precise classification. The usual caution must be taken concerning the spectral classification of Be stars, which can be influenced by the circumstellar envelope and the gravitational darkening. The rotation velocity was estimated by evaluating the first zero of the Fourier transform of several He lines. We conclude that the star rotates with  $V \sin i = 120 \text{ km s}^{-1}$ . This is a maximum value for the projected stellar velocity, because the spectral resolution is about 2 Å, which corresponds to  $\approx 120 \text{ km s}^{-1}$ .

## 5. Discussion

The two dominant frequencies found in Sect. 3 are  $f_1 = 2.465 \text{ c d}^{-1}$  and  $f_2 = 2.441 \text{ c d}^{-1}$ . Taking  $M_* = 3.8 M_\odot$  and  $R_* = 2.6 R_\odot$  for a non-flattened B 5.5V star (Huang & Gies 2006) and using the expression of Zorec et al. (2005), the critical rotation velocity is  $528 \text{ km s}^{-1}$  and  $v_{\text{crit}} = 4 \text{ c d}^{-1}$ . Frémat et al. (2005) found that typical Galactic Be stars rotate on average at 88% of their critical rate. If that is the case for CoRoT-ID 102761769,  $v_{\text{rot}} = 3.5 \text{ c d}^{-1}$  and frequencies near  $f_1$  and  $f_2$  found in this paper are much lower than the probable rotation

frequency, which implies that they are true stellar NRP pulsation frequencies. The determined  $V \sin i$  and the above critical velocity set the star inclination angle to be  $\approx 15$  degrees. From the H $\alpha$   $Ew$  (6.6 Å) and  $FWHM$  (6.2 Å), the relations given by Dachs et al. (1986) allow us to estimate the inner radius of the region emitting the bulk of the line to be  $R_\alpha \approx 4.4 R_*$ . On the other hand, the absorption cusps seen at the central region of the H $\beta$  line indicate that the inner radius of the H $\beta$  line-forming region is  $R_\beta \approx 1.95 R_*$ .

## 6. Conclusions

We have analyzed 54.6 days of photometric data of the CoRoT-ID 102761769 target ( $V = 13.08$ , B5-6 IV-Ve star) obtained during the IRa01 run. One spectroscopic observation was also performed at Roque de los Muchachos observatory (La Palma, Spain), covering the spectral region  $\lambda\lambda = 3337\text{--}8366 \text{ Å}$ . The H $\alpha$  line was observed in emission and H $\beta$  was found to be partially filled with emission. The projected rotational velocity is  $V \sin i = 120 \text{ km s}^{-1}$ , as estimated by means of the Fourier transform of HeI line profiles. The time series analysis was based on both cleanest and SSA algorithms. We found two dominant frequencies,  $f_1 = 2.465 \text{ c d}^{-1}$  and  $f_2 = 2.441 \text{ c d}^{-1}$ . A strong beating between these frequencies was uncovered when the data were reconstructed by detecting the four most significant eigenvalues found applying SSA. If the star rotates near the critical frequency, the frequencies  $f_1$  and  $f_2$  found in this paper are much lower than the probable rotation frequency.

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